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Corrigendum

Corrigendum to “Did large volcanic channel systems develop on Earth during the Hadean and Archean?”

[Precambrian Res. 246 (2014) 226–239]

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The author regrets that Kasei Valles was not depicted properly in [Fig. 1](#). The corrected figure is reprinted overleaf. The author apologises for any inconvenience caused.

DOI of original article: <http://dx.doi.org/10.1016/j.precamres.2014.03.006>.

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<http://dx.doi.org/10.1016/j.precamres.2014.11.001>

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Please cite this article in press as: Leverington, D.W., Corrigendum to “Did large volcanic channel systems develop on Earth during the Hadean and Archean?” [Precambrian Res. 246 (2014) 226–239]. Precambrian Res. (2014), <http://dx.doi.org/10.1016/j.precamres.2014.11.001>

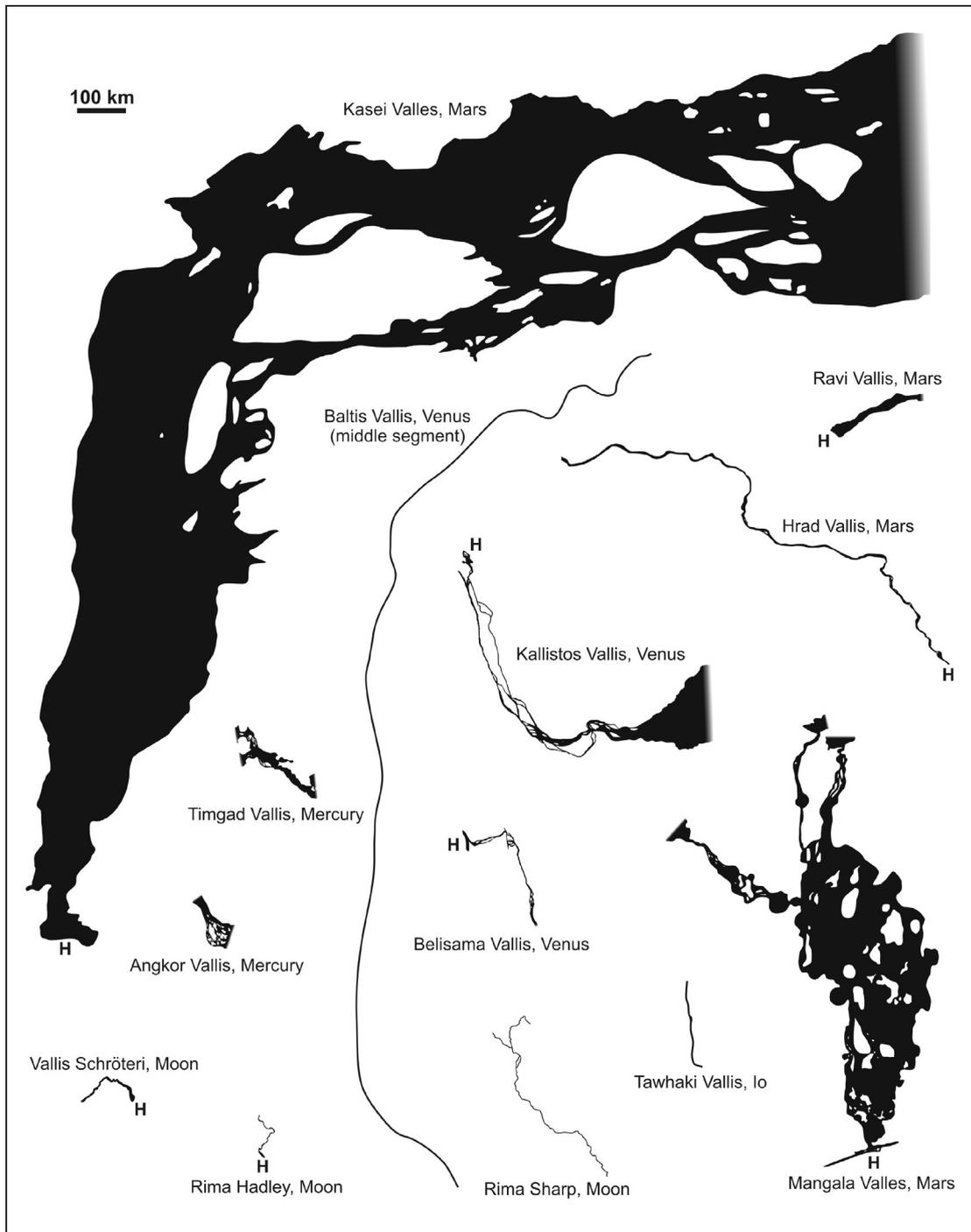


Fig. 1. Selected examples of large volcanic channels of the solar system, depicted using a common scale. The heads of systems are identified by H, except where sources are obscured or otherwise unknown. Lunar channels are typically simple and relatively small, though the longest known system (Rima Sharp) has a total length of at least 566 km (Hurwitz et al., 2013a). Some of the simplest Venusian volcanic systems (the ‘canali’) are also the longest, with Baltis Vallis remarkably possessing a length of ~6800 km despite its maximum width of ~3 km (Komatsu et al., 1993); the segment depicted here represents only ~1/4 of the total length of this system, which is considered to be the longest known channel of any kind in the solar system (Komatsu and Baker, 1996). Among the most complex Venusian channels is the Kallistos Vallis system (Baker et al., 1992). Though the longest Mercurian system has a relatively modest length of ~161 km, the widths of channel systems on this planet are as great as ~28 km (Byrne et al., 2013). The Martian outflow channels are widely interpreted as having formed by aqueous outbursts from aquifers (Wilson et al., 2009), but these systems have more recently been identified as possessing attributes that are most consistent with volcanic origins (Leverington, 2004, 2011; see also Jaeger et al., 2010; Byrne et al., 2013). The Tawhaki Vallis system of Io (Schenk and Williams, 2004) has dimensions comparable to those of relatively simple volcanic channels on bodies such as the Moon and Venus.